

Online Research @ Cardiff

This is an Open Access document downloaded from ORCA, Cardiff University's institutional repository: <https://orca.cardiff.ac.uk/id/eprint/89655/>

This is the author's version of a work that was submitted to / accepted for publication.

Citation for final published version:

Clark, Paul ORCID: <https://orcid.org/0000-0002-4834-043X> and Glover, Simon C. O. 2015. Does the CO-to-H₂ conversion factor depend on the star formation rate? Monthly Notices of the Royal Astronomical Society 452 (2) 10.1093/mnras/stv1369 file

Publishers page: <http://dx.doi.org/10.1093/mnras/stv1369>
<<http://dx.doi.org/10.1093/mnras/stv1369>>

Please note:

Changes made as a result of publishing processes such as copy-editing, formatting and page numbers may not be reflected in this version. For the definitive version of this publication, please refer to the published source. You are advised to consult the publisher's version if you wish to cite this paper.

This version is being made available in accordance with publisher policies.

See

<http://orca.cf.ac.uk/policies.html> for usage policies. Copyright and moral rights for publications made available in ORCA are retained by the copyright holders.



Does the CO-to-H₂ conversion factor depend on the star formation rate?

Paul C. Clark^{1,2★} and Simon C. O. Glover^{2★}

¹*School of Physics and Astronomy, Queen's Buildings, The Parade, Cardiff University, Cardiff, CF24 3AA, UK*

²*Zentrum für Astronomie der Universität Heidelberg, Institut für Theoretische Astrophysik, Albert-Ueberle-Str. 2, D-69120 Heidelberg, Germany*

Accepted 2015 June 16. Received 2015 June 16; in original form 2015 May 2

ABSTRACT

We present a series of numerical simulations that explore how the ‘X-factor’, X_{CO} – the conversion factor between the observed integrated CO emission and the column density of molecular hydrogen – varies with the environmental conditions in which a molecular cloud is placed. Our investigation is centred around two environmental conditions in particular: the cosmic ray ionization rate (CRIR) and the strength of the interstellar radiation field (ISRF). Since both these properties of the interstellar medium (ISM) have their origins in massive stars, we make the assumption in this paper that both the strength of the ISRF and the CRIR scale linearly with the local star formation rate (SFR). The cloud modelling in this study first involves running numerical simulations that capture the cloud dynamics, as well as the time-dependent chemistry, and ISM heating and cooling. These simulations are then post-processed with a line radiative transfer code to create synthetic ¹²CO (1–0) emission maps from which X_{CO} can be calculated. We find that for $10^4 M_{\odot}$ virialized clouds with mean density 100 cm^{-3} , X_{CO} is only weakly dependent on the local SFR, varying by a factor of a few over 2 orders of magnitude in SFR. In contrast, we find that for similar clouds but with masses of $10^5 M_{\odot}$, the X-factor will vary by an order of magnitude over the same range in SFR, implying that extragalactic star formation laws should be viewed with caution. However, for denser (10^4 cm^{-3}), supervirial clouds such as those found at the centre of the Milky Way, the X-factor is once again independent of the local SFR.

Key words: molecular processes – ISM: clouds – ISM: molecules – galaxies: ISM.

1 INTRODUCTION

In both the Milky Way, and in other local galaxies, star formation is known to take place within large clouds of molecular gas, the so-called giant molecular clouds (GMCs). Understanding the properties of these GMCs is important for the light it sheds on the process of star formation. Unfortunately, the main chemical constituent of the gas in GMCs, molecular hydrogen (H₂), is extremely difficult to observe in situ, owing to the fact that the characteristic temperature of the gas in typical GMCs (10–20 K; Bergin & Tafalla 2007) is much smaller than the temperature required to excite even the lowest rotational transition of the H₂ molecule. Therefore, most observational studies of GMCs rely on carbon monoxide (CO) as a proxy for molecular hydrogen. Observations of nearby GMCs find a surprisingly tight correlation between the integrated intensity coming from the clouds in the $J = 1 \rightarrow 0$ rotational transition line of ¹²CO and the molecular hydrogen column density of the clouds (see e.g. Dickman 1978; Sanders, Solomon & Scoville 1984; Solomon et al. 1987; Strong & Mattox 1996; Dame, Hartmann & Thaddeus

2001; or the recent review by Bolatto, Wolfire & Leroy 2013), and yield a conversion factor between CO intensity¹ and H₂ column density given approximately by (Dame et al. 2001)

$$X_{\text{CO,gal}} = \frac{N_{\text{H}_2}}{W_{\text{CO}}} = 2 \times 10^{20} \text{ cm}^{-2} \text{ K}^{-1} \text{ km}^{-1} \text{ s}, \quad (1)$$

where W_{CO} is the velocity-integrated intensity of the CO $J = 1 \rightarrow 0$ emission line, averaged over the projected area of the GMC, and N_{H_2} is the mean H₂ column density of the GMC, averaged over the same area. This latter quantity cannot be directly determined from observations of H₂ emission, but can be inferred from measurements of the total column density if the H I column density has also been mapped. The total column density of gas in the cloud can itself be determined for nearby clouds using extinction mapping (see e.g. Kainulainen et al. 2009; Froebrich & Rowles 2010; Pineda et al. 2010) or measurements of the diffuse γ -ray flux produced by interactions between high-energy cosmic rays and atomic hydrogen, atomic helium and H₂ (e.g. Digel et al. 1999).

* E-mail: paul.clark@astro.cf.ac.uk (PCC); glover@uniheidelberg.de (SCOG)

¹ Unless otherwise noted, when we refer simply to CO from this point on, we mean ¹²CO.

The fact that the values for X_{CO} derived for nearby clouds show very little variation has led to its widespread use as a general CO-to- H_2 conversion factor (commonly referred to as the ‘X-factor’), even in environments very different from the Galactic interstellar medium (ISM). It is therefore important to understand when the adoption of a constant value for X_{CO} is appropriate, and when it may be seriously misleading.

In clouds with a fixed H_2 column density, any variation in X_{CO} that occurs from cloud to cloud must be due solely to a variation in W_{CO} . Numerical simulations show that this is sensitive to three main parameters: the velocity dispersion of the gas in the cloud, the temperature of the CO-emitting gas, and the filling factor of bright CO emission (Glover & Mac Low 2011; Shetty et al. 2011a,b; Narayanan & Hopkins 2013). Increasing the gas temperature increases the typical brightness temperature of the emitting gas, thereby increasing W_{CO} and decreasing X_{CO} . However, as a change in the temperature also affects the excitation of the CO molecules, the scaling of W_{CO} with temperature is not linear, but instead is closer to $W_{\text{CO}} \propto T^{1/2}$ (Shetty et al. 2011b). Increasing the velocity dispersion also increases W_{CO} by increasing the CO linewidth, but the increased velocity dispersion can also lead to lower CO line opacities and hence lower brightness temperatures, so again the dependence is sub-linear (Shetty et al. 2011b). Finally, W_{CO} is sensitive to the fraction of the cloud traced by bright CO emission, as in cases where this is small, the *mean* integrated intensity is much smaller than the *peak* integrated intensity, owing to the effects of beam dilution. In solar metallicity clouds situated in a standard Galactic radiation field, the filling factor of the CO emission is of the order of unity (see e.g. Wolfire, Hollenbach & McKee 2010, or Section 3 below), and so small variations in its value from cloud to cloud have little effect on X_{CO} . At lower metallicities, however, the CO filling factor can become very small (Glover & Mac Low 2011; Glover & Clark 2012b) and hence X_{CO} can increase dramatically.

If we also allow the H_2 column density to vary, then this adds an additional parameter to the problem. However, in real clouds, it is unlikely that all of these parameters vary independently. In particular, we expect the filling factor of CO-bright gas to be sensitive to the mean extinction of the cloud (Glover & Mac Low 2011), which at fixed metallicity is directly related to N_{H_2} . Nevertheless, in order to understand why X_{CO} is close to constant in local GMCs, we need to understand why these four parameters do not all vary by large amounts.

If GMCs are in virial equilibrium with linewidths that satisfy the standard Galactic size–linewidth relationship (Solomon et al. 1987; Scoville et al. 1987; Roman-Duval et al. 2010)

$$\sigma \simeq 0.7 \left(\frac{R}{1 \text{ pc}} \right)^{1/2} \text{ km s}^{-1}, \quad (2)$$

where σ is the linewidth and R is a measure of the characteristic size of the cloud, then we would expect that N_{H_2} would be approximately constant from cloud to cloud (Larson 1981) and that the CO linewidth would scale only weakly with the cloud mass as $\sigma \propto M_{\text{vir}}^{1/4}$ (Bolatto et al. 2013), thereby explaining much of the constancy in X_{CO} . However, it has become clear over the last few years that this is not the only possible explanation and that virial equilibrium is not required in order to explain a constant X_{CO} . The reason is that even if GMCs are not virialized, their mean H_2 column densities are unlikely to vary by a large amount. Clouds with low values of N_{H_2} will provide insufficient dust shielding to allow widespread CO formation and hence will not be identified as ‘molecular’ clouds (Glover & Mac Low 2011; Clark & Glover 2014), while clouds with high N_{H_2} will form stars rapidly, and the resulting stellar feedback will

prevent their column densities from growing too large (Feldmann, Gnedin & Kravtsov 2012; Narayanan & Hopkins 2013).

The small variation that we find in the mean CO brightness temperature of local GMCs can also be easily understood. Since ^{12}CO emission from GMCs is usually optically thick, the observed brightness temperature is determined by the density and temperature of the gas at the point in the cloud where $\tau = 1$ (Bolatto et al. 2013). These values vary much less from cloud to cloud than do the mean values for the cloud as a whole (Molina, Glover & Federrath 2011; Narayanan & Hopkins 2013), and so there is little variation in the resulting mean brightness temperatures (Roman-Duval et al. 2010; Narayanan & Hopkins 2013).

Putting all of this together, we see that we can understand fairly well why X_{CO} is approximately constant for local GMCs. However, the behaviour of X_{CO} as we move to environments with much stronger UV radiation fields and higher cosmic ray fluxes is less easy to predict. We would expect that as we increase the strength of the ambient radiation field, more photodissociation of CO will occur. This will decrease the filling factor of the CO emission and may also decrease the CO linewidth if the CO-emitting gas is no longer well distributed throughout the volume of the cloud. These effects will tend to decrease W_{CO} and hence increase X_{CO} . At the same time, we also expect a stronger UV field to lead to greater dissociation of H_2 , and hence a decrease in N_{H_2} . In addition, we would expect the cloud to be systematically warmer, owing to the higher photoelectric heating rate in low A_V regions, and the greater cosmic ray heating rate in high A_V regions. This means that the mean brightness temperature of the CO emission may also be larger. Together, these two effects will tend to decrease X_{CO} . It is not immediately clear which set of effects will dominate, and hence whether one would expect X_{CO} to increase or decrease as the strength of the UV radiation field and the cosmic ray flux increase.

In this paper, we investigate this issue with the help of numerical simulations. We simulate the coupled chemical, thermal and dynamical evolution of several model clouds in a variety of different environments, spanning a range of 100 in terms of UV field strength and cosmic ray ionization rate (CRIR). We model and analyse the resulting emission in the 1–0 line of ^{12}CO , and thereby are able to explore how X_{CO} varies as a function of the external environment in clouds with a fixed mass and mean volume density. In Section 2, we briefly describe our numerical approach, and in Section 3, we explore how the changing environment affects the physical structure of the clouds. In Section 4, we examine how X_{CO} varies from cloud to cloud, and how this depends on the definition used for the cloud boundary. We discuss some implications of our results in Section 5 and conclude in Section 6.

2 NUMERICAL METHOD

2.1 Details of the algorithms

The computations presented in this paper were performed using a modified version of the publicly available smoothed particle hydrodynamics (SPH) code GADGET-2 (Springel 2005). A major modification to the original GADGET-2 is the inclusion of time-dependent chemistry. The version we use here follows the formation and destruction of H_2 as introduced by Glover & Mac Low (2007a,b), in addition to the simplified treatment of the CO chemistry that was proposed by Nelson & Langer (1999). In Glover & Clark (2012a), it was shown that this CO network has a similar accuracy to the more exhaustive treatment of Glover et al. (2010), but incurs around one third of the computational cost. We do not include freeze-out of

CO on to dust grains in this model, but as this generally occurs only in regions where the ^{12}CO emission is already optically thick (Goldsmith 2001), we do not expect this omission to significantly affect our results.

As well as following the chemical evolution of the gas, we also model its thermal evolution. We account for dynamical heating due to shocks and adiabatic compression and cooling due to adiabatic rarefaction in the same fashion as in the unmodified version of GADGET-2. In addition, we also account for the main radiative and chemical heating and cooling processes occurring in the ISM. These include fine structure cooling from C^+ , C and O, molecular line cooling from H_2 and CO, photoelectric heating and cosmic ray heating. At high gas densities, collisions between gas particles and dust grains also play an important role in regulating the thermal energy balance, cooling the gas if $T_{\text{gas}} > T_{\text{dust}}$ and heating it if $T_{\text{dust}} > T_{\text{gas}}$. Full details of how we treat these processes and a number of other, less important, contributors to the overall thermal energy balance can be found in our previous papers (Glover & Mac Low 2007a,b; Glover et al. 2010; Glover & Clark 2012a), and a summary of the most important processes included can be found in fig. 4 in Clark et al. (2013).

The attenuation of the interstellar radiation field (ISRF) is treated using the TREECOL algorithm, introduced by Clark, Glover & Klessen (2012a). In this paper, the spectral shape of the ISRF is based on the prescription of Draine (1978) in the ultraviolet and Black (1994) at longer wavelengths. The strength of the ISRF is varied in the different simulations as described in Section 2.3 below. The clouds are assumed to be bathed in a uniform ISRF, and TREECOL is used to compute the attenuated spectrum that reaches each SPH particle in the computational volume.

Since the above-mentioned papers, we have made two significant changes to the chemical model. The first was to update our treatment of the photodissociation of CO from the prescription given in Lee et al. (1996), to that described in the recent paper of Visser, van Dishoeck & Black (2009). The second change that we have made to the chemical model is the inclusion of the effects of cosmic ray ionization of atomic carbon



and cosmic ray induced photodissociation of C and CO



These processes were not included in the chemical model described in Glover & Clark (2012a). We assume that the rates of all three processes are proportional to ζ_{H} , the CRIR of atomic hydrogen. For processes 3 and 4, we use the rates given in Woodall et al. (2007) as a basis, but rescale them by a factor $\zeta_{\text{H}}/\zeta_{\text{H}, \text{W07}}$, where $\zeta_{\text{H}, \text{W07}}$ is the value of the CRIR adopted by Woodall et al. (2007). Similarly, to compute the rate of process 5, we use the value given in Gredel, Lepp & Dalgarno (1987) as a basis, but rescale it to make it consistent with our choice of ζ_{H} .

The ^{12}CO (1–0) emission maps that form the basis of the analysis in this paper were created using the RADMC-3D radiative transfer code.² The level populations in the non-local thermodynamic equilibrium limit are computed using the large velocity gradient approximation (Sobolev 1957), as implemented in RADMC-3D by Shetty et al. (2011a,b). We assume that collisions with H_2 dominate

the excitation of the CO rotational levels and adopt the values for the collisional excitation rate coefficients given in the Leiden Atomic and Molecular data base (Schöier et al. 2005).

2.2 Cloud properties

All of the clouds in this study start as uniform spheres of gas, on to which a three-dimensional turbulent velocity field has been superimposed. Two cloud masses are examined in this paper: a ‘low-mass’ cloud, with a mass of $10^4 M_{\odot}$, and a ‘high-mass’ cloud of mass $10^5 M_{\odot}$. In most of the runs, we take the initial hydrogen nuclei number density of the gas to be $n_0 = 100 \text{ cm}^{-3}$, yielding initial radii for the low- and high-mass clouds of 8.8 and 19 pc, respectively. In addition, we performed two simulations of high-mass clouds with a much higher initial density, $n_0 = 10^4 \text{ cm}^{-3}$; these denser clouds had an initial radius of around 4 pc.

The mass resolution in this study is kept fixed, with the mass of an SPH particle being set at $0.005 M_{\odot}$. The low-mass clouds therefore have 2×10^6 SPH particles, while the high-mass clouds have 2×10^7 SPH particles. The minimum resolvable self-gravitating mass element in this calculation is therefore $0.5 M_{\odot}$ in both cases (Bate & Burkert 1997). In practice, this means that we can follow the collapse of the cloud until the gas number density reaches a value of around 10^6 cm^{-3} . At this point, we halt the collapse and perform the analysis that is presented in this paper. We note that we would not expect the results presented here to change if we were to follow the collapse to higher densities, as at these densities, the CO (1–0) line will be highly optically thick (see e.g. Goldsmith 2001), and in any case much of the CO will be frozen out on to dust grains.

The velocity fields are generated with a ‘natural’ mix of solenoidal and compressive modes (i.e. a ratio of 2:1). This is generated on a 128^3 grid. This velocity field is left to freely decay in shocks, rather than being continuously driven during the course of the simulation. In most of the simulations, the initial energy in the velocity field, E_{kin} , was set to be half the initial gravitational potential energy of the cloud, E_{grav} , so that the clouds are initially in virial equilibrium. This means that the initial rms turbulent velocity v_{rms} is 2.4 km s^{-1} for the low-mass clouds and 5.2 km s^{-1} for the high-mass clouds. In terms of the virial parameter

$$\alpha_{\text{vir}} = \frac{E_{\text{kin}}}{E_{\text{grav}}}, \quad (6)$$

these clouds have $\alpha_{\text{vir}} = 0.5$. In addition, we also examine one high-density cloud for which we set $\alpha_{\text{vir}} = 2$, so that the cloud is initially unbound (see Section 4.3). Note that most of the simulations presented here adopt the same random seed for the turbulent velocity field. However, we also run one of the simulations with a different seed, to gauge the sensitivity of our results to the underlying cloud structure that is created by the turbulent motions.

From previous modelling, it has been found that in cold gas with densities of around 100 cm^{-3} , most of the hydrogen is in the form of H_2 . In contrast, the carbon in this gas is still predominantly in the form of C^+ (see e.g. Clark et al. 2012b; Smith et al. 2014). The initial chemical state of the gas in the majority of the runs presented in this study is motivated by this previous work. We start most of our runs with all of the hydrogen already in the form of H_2 , but assume that the carbon and oxygen are present in the form of C^+ and O, respectively. We adopt total abundances (relative to the number density of hydrogen nuclei) for the carbon and oxygen nuclei of $x_{\text{C}} = 1.4 \times 10^{-4}$ and $x_{\text{O}} = 3.2 \times 10^{-4}$, respectively, consistent with the values measured in the local ISM (Sembach et al. 2000).

² <http://www.ita.uni-heidelberg.de/~dullemond/software/radmc-3d/>

Similarly, we adopt a dust-to-gas mass ratio of 0.01, consistent with the value in the local ISM.

In our runs with very high UV fields, it is unclear whether starting with all of the hydrogen in molecular form is a good approximation, as in this case, one would expect the equilibrium H_2 fraction in low-density gas to be much smaller than in the models of Clark et al. (2012b) and Smith et al. (2014). To address this uncertainty, we ran two additional high-mass models with the hydrogen initially in atomic form.

The post-processing of the simulation data in RADMC-3D first requires that the SPH particle data is interpolated on to a regular Cartesian grid. This is done using the standard SPH smoothing formalism. To ensure that we catch small, high-density pockets of gas, we employ a RADMC-3D cell size of 0.068 pc, such that the high-mass cloud calculations have 656^3 cells and the low-mass cloud calculations have 256^3 cells. We find that such a resolution is sufficient to get converged values for the probability density function of CO (1–0) emission, from which it follows that the various mean values that we examine later are also converged. Once the position–position–velocity maps are obtained, we then integrate the emission along the z -axis to create maps of the integrated intensity, $W_{\text{CO}(1-0)}$.

2.3 The ‘star formation rate’

In this study, we vary two of the environmental conditions that can affect the chemical balance of clouds: the strength of the ISRF and the CRIR. Both of these are thought to vary with the local star formation rate (SFR), and so we assume here that the strength of these processes can be used as a proxy for the SFR.

As mentioned above, the ISRF used here is taken to have a shape described by a combination of the Black (1994) and Draine (1978) radiation fields. In one set of runs – those representing clouds in an environment with an SFR similar to that in the local ISM, which we denote as SFR_0 – we adopt the same normalization for the ISRF as in the papers of Draine (1978) and Black (1994). Our fiducial ISRF therefore has a strength $G_0 = 1.7$ in Habing (1968) units.

For this study, we are mainly interested in how the ISRF heats the gas and affects its chemical state. Photons with energies above 6 eV are responsible for the photoelectric heating (the dominant heat source in low extinction regions of these clouds), and photons with energies above 11.2 and 11.5 eV are responsible for dissociating H_2 and CO, respectively. Since most of the photons in this part of the ISRF come from massive, young stars, it is reasonable to assume that, to a first approximation, the strength of the relevant portion of the ISRF scales linearly with the local SFR. At longer wavelengths, the ISRF is dominated by older stellar populations and this assumption is less well founded. However, the strength of the ISRF at these wavelengths has little effect on the temperature or chemistry of the gas, and so for simplicity, we assume that in regions with higher SFRs, we can simply scale the entire ISRF upwards, rather than changing its spectral shape.

We also assume that the CRIR scales linearly with the SFR. This assumption is reasonable: supernova remnants are the main source of cosmic rays in the Galactic ISM, and the lifetime of a typical cosmic ray within the Galaxy is around 15 Myr (Ferriere 2001), so the cosmic ray energy density, and hence the CRIR, should track the SFR fairly closely.

In this study, in addition to the runs representing the behaviour of clouds in the local ISM, with $\text{SFR} = \text{SFR}_0$, we also perform simulations where we increase the strength of the ISRF and the CRIR by factors of 10 or 100, corresponding to SFRs of 10 or 100 SFR_0 , respectively.

3 GENERAL CLOUD STRUCTURE

In the left-hand columns of Figs 1 and 2, we show column density images of the gas, taken at a point in the simulations just before the onset of star formation in our fiducial clouds – i.e. those seeded with our standard turbulent velocity field, and which started life with all their hydrogen in the form of H_2 . This particular stage in the cloud’s evolution is when we perform the analysis that forms the basis of this paper. For most of our simulated clouds, this corresponds to a time $t \sim 1\text{--}2$ Myr after the beginning of the simulation, as shown in Table 1.

In both the 10^4 and $10^5 M_\odot$ clouds, we see a similar change in behaviour of the strength of the ISRF and CRIR are increased to mimic progressively higher rates of ambient star formation activity: the filamentary structure imposed by the cloud’s turbulence and self-gravity is less pronounced as we move to higher SFRs. This is simply a result of gas temperatures rising due to the combined effects of the enhanced photoelectric emission and CRIRs. This systematic increase in the temperature as we increase the SFR is also illustrated in Fig. 3, where we show temperature–density phase diagrams for four of our runs, colour coded by the CO abundance.

In addition, we see from Figs 1 and 2 that the overall morphology of the clouds changes as the SFR increases. In this respect, there are two trends. The first is that the dense, filamentary structure becomes confined to a progressively smaller region as the SFR increases. This makes sense, as the gas is more easily structured near the centre of the cloud where the extinction is higher and the heating effects of the photoelectric emission are less. As the SFR increases, the heating by photoelectric emission becomes progressively more important in shaping the gas, and so the regions structured by the turbulence and gravity are those towards the centre.

The second effect is that the structure of the cloud envelope changes. The SPH particles have open (i.e. vacuum) boundary conditions, and so the particles are free to expand into the surrounding space. The stronger the ISRF, the more the surface of the cloud is heated, causing it to ‘boil off’ into the void, creating the halo that we can see around the clouds. In the case of the 100 SFR_0 clouds (i.e. those illuminated by radiation fields and cosmic ray fluxes that are 100 times greater than the values in the local ISM; see Section 2.3), the ISRF is so strong that this halo becomes hot enough to push back on the cloud, creating the sharp boundary that we see in Figs 1 and 2. Although dramatic, this feature of the increasing SFR is actually of little interest to this current study, as both the H_2 and CO lie within the cloud boundary.

The middle panels in Figs 1 and 2 show the H_2 column densities images at the same point in the cloud evolution. In general, we see that the H_2 column density tracks the total column density extremely well for the bulk of the cloud. Given that all the clouds in these two figures were initialized with all their hydrogen in molecular form, this is not surprising.

The main difference between the left-hand and middle panels in Figs 1 and 2 is that the H_2 in the outer envelope has been photodissociated by the ISRF, and so does not trace the very low density (both column and volume) that surrounds the main cloud structure. We also see that the column density at which H_2 starts to appear increases with increasing SFR, moving from around 10^{20} to around 10^{21} cm^{-2} as the SFR increases from the solar neighbourhood value to a factor of 100 greater. This behaviour is consistent with the predictions of detailed models for the structure of photodissociation regions (see e.g. Krumholz, McKee & Tumlinson 2008; Sternberg et al. 2014). The fact that a large change in the strength of the ISRF leads to a relatively small change in the

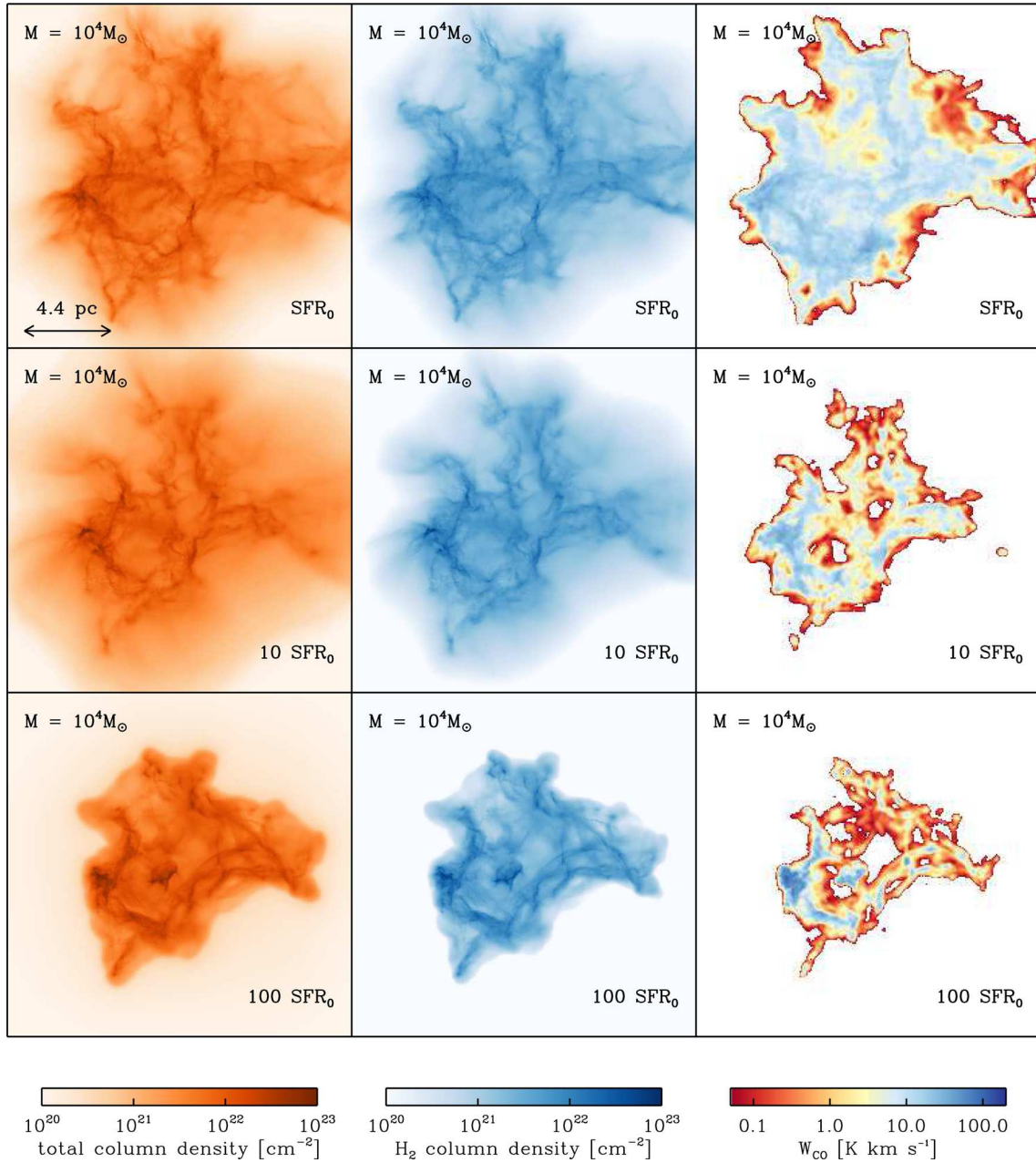


Figure 1. The columns show the total column density (left), the H_2 column density (middle) and the velocity-integrated CO intensity in the 1–0 line (right) for the ‘standard’ SFR_0 (top row), $10 \times \text{SFR}_0$ (middle row) and $100 \times \text{SFR}_0$ (bottom row) in the $10^4 M_\odot$ clouds. All of the images were made at the point at which the density exceeded 10^6 cm^{-3} in the first collapsing core in each simulation.

column density of gas required to effectively shield H_2 is a consequence of the fact that dust shielding is an exponential function of the column density, which implies that in the dust-dominated region, the critical column density scales only logarithmically with the SFR.

The integrated emission in the 1–0 line of ^{12}CO , W_{CO} , is shown in the right-hand column in Figs 1 and 2, and shows the most striking variation with increasing SFR. For the SFR_0 clouds, we see that the surface filling factor of the bright CO emission is very similar to that of the H_2 gas. Despite this, it is not a good tracer of the physical structure of the cloud: some features seen in the emission map are far more blobby than their H_2 counterparts and others are missing

entirely. This behaviour is expected (see e.g. Ballesteros-Paredes & Mac Low 2002; Goodman, Pineda & Schnee 2009; Shetty et al. 2011a,b; Beaumont et al. 2013) and is simply a result of the large optical depth of the 1–0 line – the features that we see are the $\tau \sim 1$ surfaces, which are a function of CO column density, velocity and temperature, and so often bear little resemblance to the structure we see in the H_2 column density images.

As we move to higher SFRs, we see that the surface filling factor of the CO emission progressively decreases, a feature we see in all of the cloud models run in this study. In the case of the 100 SFR_0 clouds, the effect is extreme, with emission falling below 1 K km s^{-1} for much of the cloud’s interior – the sensitivity limit in many

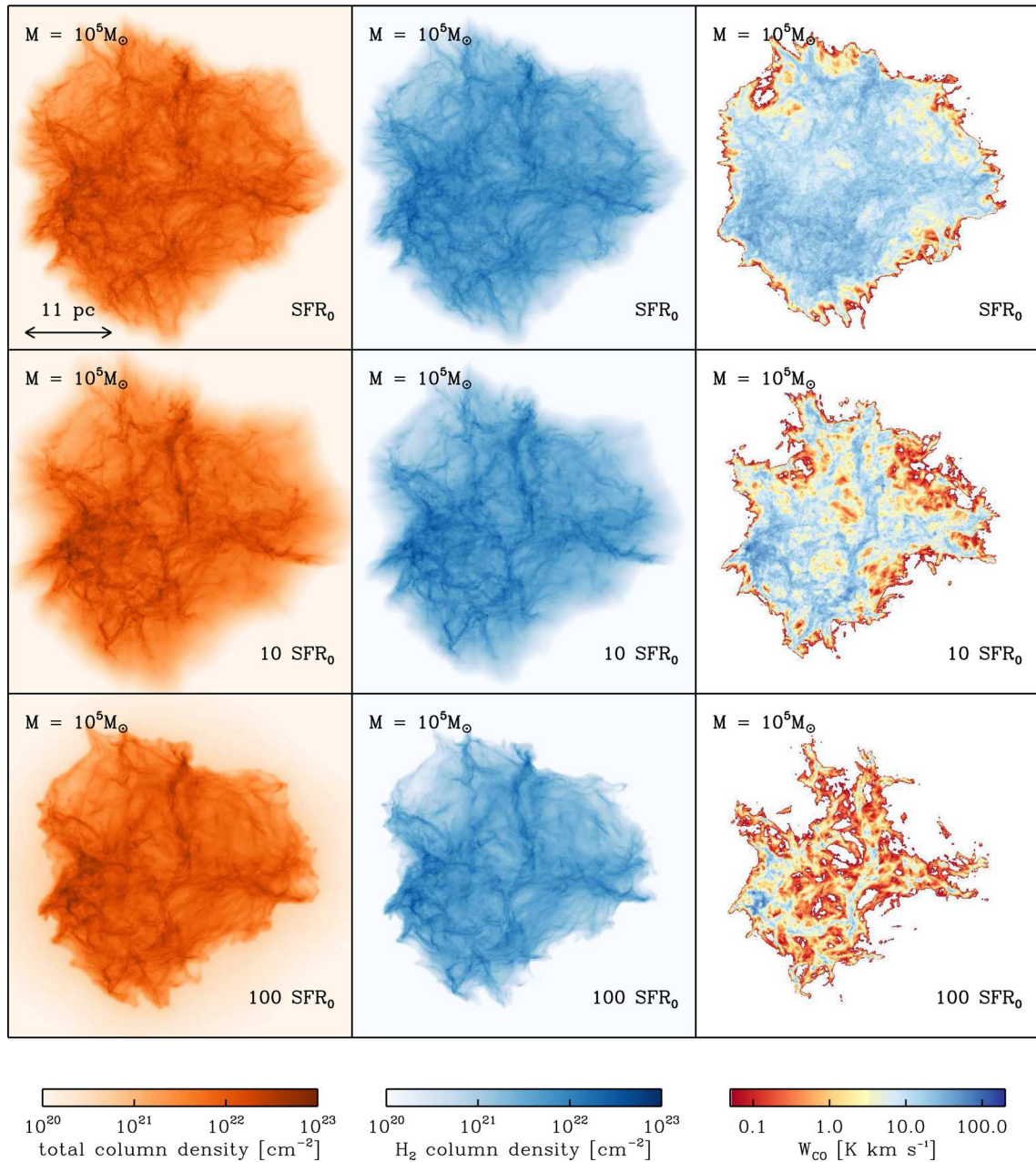


Figure 2. As Fig. 1, but for the $10^5 M_{\odot}$ clouds. Note the change of physical spatial scale, as illustrated in the top left panel.

large-scale surveys – and even below 0.1 K km s^{-1} along lines of sight with H_2 column densities as high as a few 10^{21} cm^{-2} .

However, in the cases with the higher SFR, we see that the *peak* CO emission becomes brighter towards regions of higher column density. This is a consequence of increased heating in the cloud, which results in a higher overall excitation temperature of the CO. The broader linewidth in the gas also helps to the lower the optical depth, allowing more of the emission from the line to escape (Shetty et al. 2011a,b). Due to the higher spatial resolution in the reproduced image (the true resolution is the same in all images), this effect is perhaps most easily seen in the $10^4 M_{\odot}$ clouds in Fig. 1. Here, we see the brightening of the clump towards the left of the $W_{\text{CO}(1-0)}$ image, which goes from having a peak brightness of around 30 K km s^{-1} for SFR_0 , to nearly 80 K km s^{-1} in the case of 100 SFR_0 .

4 X-FACTOR VARIATIONS

4.1 Computing the X-factor

In the previous section, we saw that as the SFR in the ambient cloud environment is increased, the CO emission traces progressively less of the H_2 distribution. However, we have also seen that the peak emission in the cloud is higher in the case of high SFR. Bolatto et al. (2013) have suggested that these two effects may effectively cancel out, resulting in an X-factor that does not depend on the cloud's environment. In this scenario, the greater emission that comes from the dense bright peaks in the high SFR case makes up for the fact that much of the cloud is not emitting in CO at all, since the molecule is easily photodissociated. Using our self-consistent models for N_{H_2} and W_{CO} , we can now explore this idea.

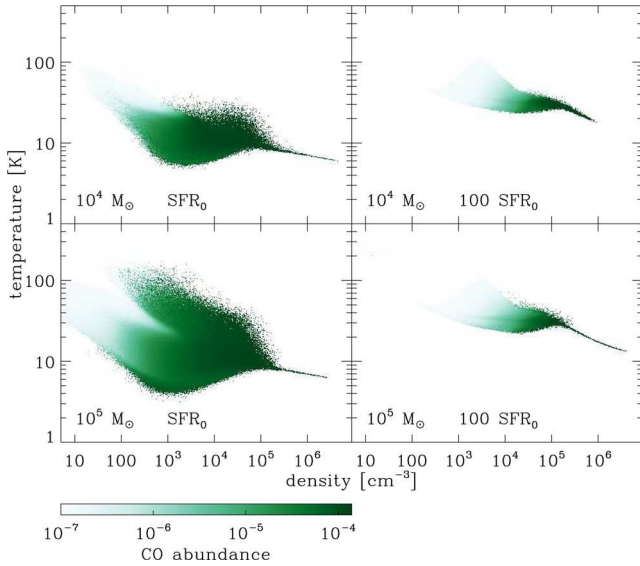


Figure 3. Phase diagrams for four of our simulated clouds, illustrating the temperature structure of the gas at the end of the simulation. The diagrams are colour coded using the mean CO abundance at each point in temperature–density phase space.

In what follows, we compare how X_{CO} varies as we alter the environmental conditions in our suite of turbulent clouds. We calculate X_{CO} for each cloud via

$$X_{\text{CO}} = \frac{\langle N_{\text{H}_2} \rangle}{\langle W_{\text{CO}} \rangle}, \quad (7)$$

where the averages are taken over a specified set of the pixels in the images from Figs 1 and 2.

In our evaluation of equation (7), we take two approaches. First, we examine the values of X_{CO} obtained when we average over all of the pixels in the images. In our second approach, we instead restrict the calculation to include only those pixels with CO 1–0 integrated intensities above some threshold value, $W_{\text{CO, cut}}$. These two approaches mimic different types of observations. The first approach is more appropriate when considering unresolved molecular clouds that could inhabit the beam in an extragalactic CO observation. In this case, the CO emission may not trace all of the H_2 gas in the beam, i.e. there is a ‘dark’ molecular component surrounding the CO-bright portion of the cloud, as recently highlighted by Wolfire

et al. (2010). The second approach, in which we only examine pixels with emission above $W_{\text{CO, cut}}$, is more applicable to studies of nearby, resolved molecular clouds, where it is common to define the extent of a cloud in terms of its CO emission, a definition which inevitably leads to one ignoring those regions without detectable CO emission (see e.g. Pineda, Caselli & Goodman 2008; Lee et al. 2014).

The mean values of X_{CO} that we derive from the simulations, and how they scale with the SFR, are shown in Fig. 4. The top two panels show the results that we obtain for two different CO detection thresholds, $W_{\text{CO, cut}} = 3 \text{ K km s}^{-1}$ and $W_{\text{CO, cut}} = 0.3 \text{ K km s}^{-1}$, while the bottom panel shows the results in the case where we do not apply a detection threshold. We also indicate the canonical local value of the X-factor, $X_{\text{CO, gal}}$ (the horizontal dashed line) and illustrate the typical scatter in the value of X_{CO} derived for the discs of nearby spiral galaxies (the shaded region), as summarized in Bolatto et al. (2013).

4.2 Clouds with $\text{SFR} = \text{SFR}_0$

We begin our study of the effects of the local environment on X_{CO} by looking at the behaviour of clouds located in an environment with an SFR similar to the local ISM, i.e. with $\text{SFR} = \text{SFR}_0$ (see Section 2.3). The first point to note is that when we consider clouds in this environment that have properties similar to those of local molecular clouds (i.e. virial parameters close to unity and mean initial densities around 100 cm^{-3}), we find that X_{CO} does not appear to be sensitive to mass. Both our low-mass and our high-mass clouds yield a mean X-factor of between 4 and $5 \times 10^{20} \text{ cm}^{-2} \text{ K}^{-1} \text{ km}^{-1} \text{ s}$, consistent with the large body of literature suggesting that the X-factor is not strongly dependent on the cloud mass (see e.g. Solomon et al. 1987; Blitz et al. 2007; Hughes et al. 2010; Bolatto et al. 2013). Also, we find that the X-factor is relatively insensitive to the value of $W_{\text{CO, cut}}$, indicating that the CO emission is a good proxy for the total H_2 present in such clouds, even though the emission is clearly not tracing all of the molecular gas.

While this is encouraging, it should also be noted that the values we recover for X_{CO} in this case are roughly twice as large as the canonical Galactic value, $X_{\text{CO, gal}}$, and on the edge of the range of values found for the discs of nearby spiral galaxies (Bolatto et al. 2013). The reason for this discrepancy is not completely clear, but there are several factors that may contribute to our recovering a systematically higher X_{CO} . The first of these stems from the way

Table 1. List of simulations. All number densities are given with respect to the number of hydrogen nuclei, since this is invariant of the chemical state of the gas. The definition of α_{vir} can be found in equation (6), while SFR_0 is defined in Section 2.3. The property t_{end} denotes the point at which the simulation is stopped, and the cloud details are fed to the line radiative transfer code. This is also the time at which the first star forms in the simulation.

Mass (M_{\odot})	n_0 (cm^{-3})	α_{vir}	SFR (SFR_0)	Notes	t_{end} (Myr)
10^4	100	0.5	1		1.83
10^4	100	0.5	10		2.09
10^4	100	0.5	100		1.91
10^4	100	0.5	100	Different seed	2.18
10^5	100	0.5	1		1.17
10^5	100	0.5	10		1.52
10^5	100	0.5	100		1.39
10^5	100	0.5	1	Atomic ICs	1.31
10^5	100	0.5	100	Atomic ICs	1.26
10^5	10^4	0.5	100	Galactic Centre style cloud	0.1
10^5	10^4	2	100	Galactic Centre style cloud	0.1

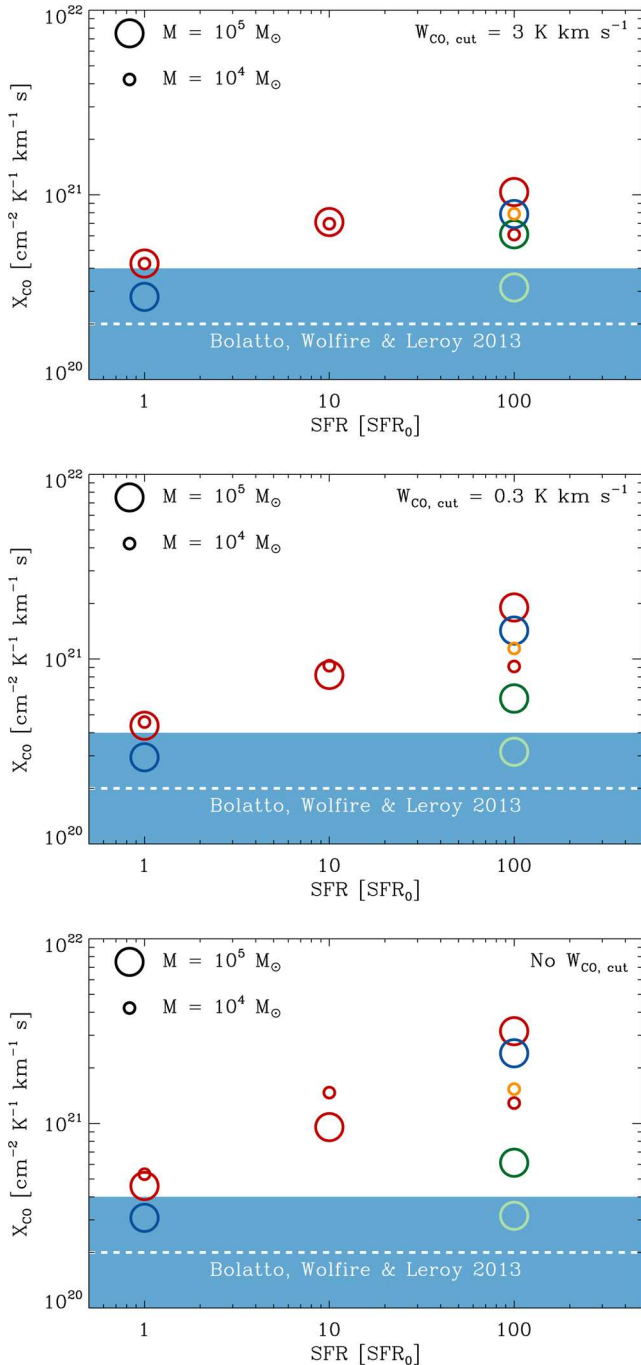


Figure 4. Mean CO-to-H₂ conversion factor, X_{CO} , in each of our clouds, as a function of our SFR proxy (see Section 2.3 for details). Red and blue circles denote the results from the $n_0 = 100 \text{ cm}^{-3}$ clouds with fully molecular and fully atomic initial conditions, respectively. The orange circles correspond to a cloud with $n_0 = 100 \text{ cm}^{-3}$ run using a different random seed for the turbulent velocity field. The dark and light green circles denote the $n_0 = 10^4 \text{ cm}^{-3}$ clouds with, respectively, $\alpha_{\text{vir}} = 0.5$ and 2. In the top two panels, we compute X_{CO} using integrated intensities and H₂ column densities only for those lines of sight with an integrated intensity above the specified threshold, $W_{\text{CO, cut}}$. In the bottom panel, we show the results that we obtain if we do not apply an integrated intensity threshold. In each case, we compute X_{CO} by dividing the mean H₂ column density in the considered area by the mean value of W_{CO} from the same area. The horizontal dashed line shows the canonical local value of X_{CO} , and the shaded area indicates the typical scatter in the value of X_{CO} derived observationally for the Milky Way and other nearby spiral galaxies (see Bolatto et al. 2013).

in which we treat the effects of self-shielding in our simulations. When calculating each particle’s sky maps of H₂ and CO column densities, the `TREECOL` algorithm includes contributions from all the molecules present in the cloud, regardless of the velocity at which they are moving with respect to the target particle. In reality, much of the cloud will not contribute to the self-shielding of a given location, since the cloud’s supersonic motions will Doppler shift the molecules out of the line-profile. Neglecting this effect means that we overestimate the effectiveness of H₂ and CO self-shielding, and hence overproduce the abundances of both molecules. We expect the error to be larger for H₂ than for CO, since self-shielding plays a much greater role in determining the H₂ abundance than the CO abundance, and therefore it will lead to a systematic increase in our values of X_{CO} compared to those for real clouds. The importance of this effect will depend on the size of the velocity dispersion in the cloud, but should not depend strongly on our choice of SFR, and it therefore represents a systematic error that biases all of our derived values of X_{CO} to slightly higher values.

Another factor that may contribute to the discrepancy between our results and $X_{\text{CO, gal}}$ is our choice to start the simulations with the hydrogen in fully molecular form. Although this approximates the behaviour that we find in previous studies of cloud formation (Clark et al. 2012b; Smith et al. 2014), in practice the hydrogen in the assembling clouds is never *completely* molecular – there is always some atomic component associated with the cloud. In equilibrium, we expect the surface density of atomic hydrogen associated with the clouds to be of the order of $10 \text{ M}_{\odot} \text{ pc}^{-2}$ (Krumholz et al. 2008; Wolfire et al. 2010; Sternberg et al. 2014), but the time taken to reach this equilibrium can be longer than the lifetime of the cloud. In our simulations with fully molecular initial conditions, we will therefore overestimate the H₂ content of the clouds by around 10–20 per cent, and hence will overestimate X_{CO} by a similar factor. In order to verify that the effect on X_{CO} is not greater than this, we ran two simulations of our high-mass cloud in which we started with all of the hydrogen in atomic form. In one of these simulations, we adopted the solar neighbourhood value for SFR (i.e. $\text{SFR} = \text{SFR}_0$), while in the other we set $\text{SFR} = 100 \text{ SFR}_0$. The values of X_{CO} that we derived from these two runs are indicated in Fig. 4 by the blue circles. We see that, as expected, the use of atomic rather than molecular initial conditions results in a decrease in X_{CO} , but that the effect is not large and seems to be independent of our choice of SFR.

Finally, our decision to focus on clouds with a virial parameter $\alpha_{\text{vir}} = 0.5$ may also lead to us deriving a larger value of X_{CO} than that found for large samples of Galactic GMCs. As discussed in the introduction, we know that the mean integrated CO intensity of a GMC depends sensitively on the CO linewidth for that GMC (see e.g. Shetty et al. 2011a,b). Increasing the linewidth while keeping all other cloud properties the same leads to an increase in W_{CO} and hence a decrease in X_{CO} . Similarly, decreasing the linewidth decreases W_{CO} and increases X_{CO} . This implies that X_{CO} should depend to some extent on α_{vir} , in the sense that clouds with larger virial parameters that survive for long enough to form CO should have smaller values of X_{CO} than similar clouds with smaller values of α_{vir} . In any large sample of GMCs, we would not expect to find many clouds with $\alpha_{\text{vir}} \ll 1$: any clouds that do initially have $E_{\text{kin}} \ll E_{\text{grav}}$ will undergo gravitational collapse, inducing motions that rapidly increase E_{kin} driving α_{vir} towards 1 if the clouds continue to collapse or 0.5 if they settle into virial equilibrium (Ballesteros-Paredes et al. 2011). On the other hand, there is both theoretical (Dobbs, Burkert & Pringle 2011; Ward, Wadsley & Sills 2014) and observational (Heyer et al. 2009) evidence that many GMCs are

gravitationally unbound, with $\alpha_{\text{vir}} > 1$. It is therefore plausible that the mean value of α_{vir} for the population of GMCs in the Milky Way is of the order of unity or larger. If so, then it follows that the clouds that we model in this study have systematically smaller CO linewidths than the average Galactic GMC, helping to explain why the values of X_{CO} that we derive for them are larger than $X_{\text{CO, gal}}$.

4.3 Dependence on the SFR

We now come to the question at the heart of this paper: does the value of X_{CO} vary with the local ambient SFR? While the results displayed in Fig. 4 clearly show that X_{CO} is indeed a function of SFR, the extent to which the X-factor varies depends on several parameters, and we will now discuss these in turn.

First, examining the red circles in Fig. 4, we see that, in general, the mean X-factor derived from the cloud maps increases as SFR increases. Given that CO is more easily dissociated than H_2 , this makes sense: the CO is doing a progressively poorer job of tracing the true extent of the molecular gas as the ambient SFR increases. However, we see that the effect is more pronounced for the high-mass clouds than for the low-mass clouds. Indeed, there is a slight decrease in the X-factor of the 100 SFR_0 case with respect to the 10 SFR_0 case for the low-mass clouds. So while the X-factor at solar neighbourhood values of the SFR (i.e. SFR_0) is independent of mass, this is not true when the SFR is increased.

The reason behind this behaviour in the low-mass clouds can be inferred from Fig. 5, where we show the fraction of the cloud's total CO emission that comes from pixels below a given W_{CO} . For progressively higher SFR, we see that fraction of the overall integrated emission coming from very bright lines of sight depends on the cloud mass. For the low-mass clouds at high SFR, the majority of the emission comes from regions with integrated intensities of 10 K km s^{-1} or greater, while the opposite is true for the high-mass clouds, in that they gain most of their emission from lines of sight with lower W_{CO} . The idea summarized by Bolatto et al. (2013) that a few bright lines of sight can compensate for the lack of CO elsewhere in the cloud only appears to hold for our low-mass clouds.

The reason why the clouds behave differently has to do with their column density distributions. In the low-mass clouds, the mean column density is low enough that CO is efficiently photodissociated throughout much of their volume as the SFR is increased. The CO that survives is confined to the highly overdense regions produced by turbulent compression and gravitational collapse, in which the shielding is more effective. Because of the high density, the excitation temperature of the CO in these regions is high, and so is the integrated intensity of the CO emission. These clouds therefore have many regions with very low W_{CO} and a few with very high W_{CO} that dominate the total emission. In the high-mass clouds, on the other hand, the mean column density of the gas is larger and the CO is therefore better able to resist photodissociation. Because of this, CO survives in lower density regions in these clouds even when the SFR is increased significantly. The excitation temperature of the CO in these regions is lower than in the dense cores, and so the integrated CO intensities coming from the lower density gas are smaller, but the much larger filling factor of these regions compensates for this, and leads to their contribution dominating the total emission from the cloud.

It is plausible that molecular clouds in environments with higher SFRs will be denser, on account of the higher ambient gas pressure in these environments (see e.g. Elmegreen, Klessen & Wilson 2008; Colombo et al. 2014). Some support for this picture comes from observations of molecular clouds in the Central Molecular Zone of

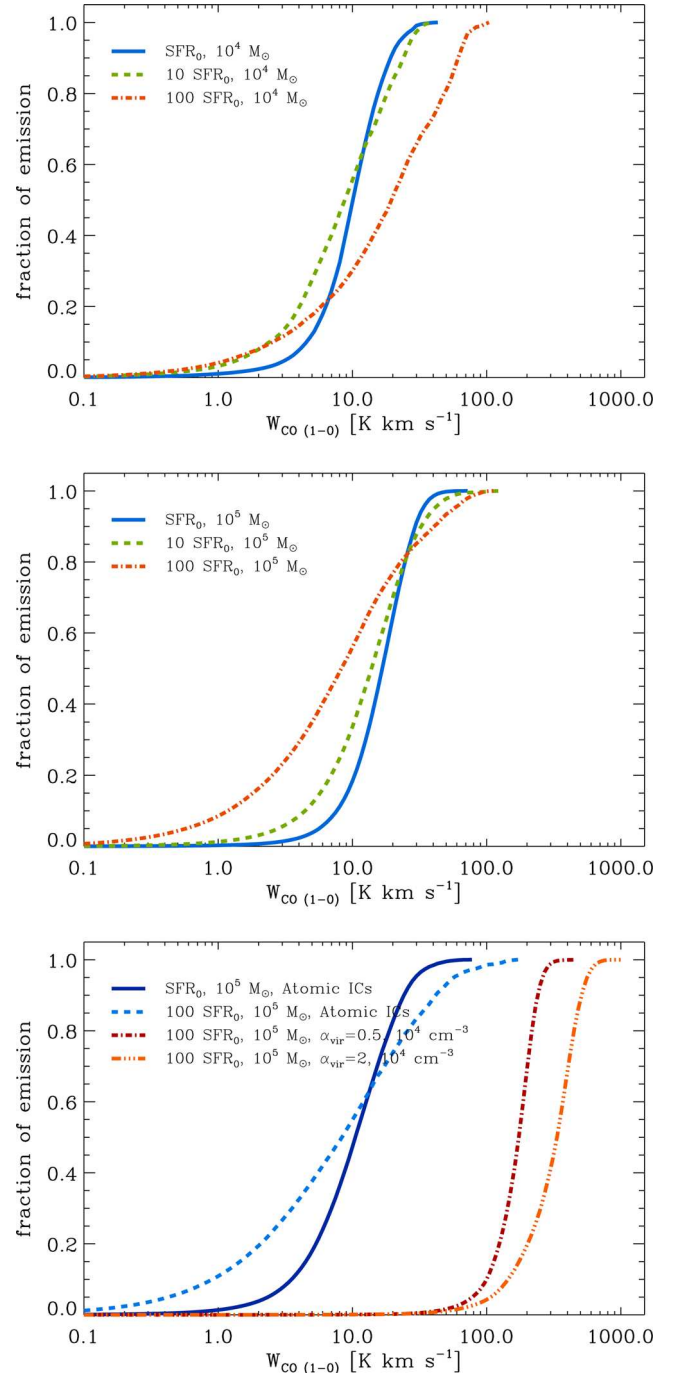


Figure 5. Fraction of the total emission that arises from pixels of a given W_{CO} and below.

the Milky Way. The SFR in this region is roughly 50–100 times that in the solar neighbourhood (Bonatto & Bica 2011; Longmore et al. 2013a) and the molecular clouds located there have high mean densities, of the order of 10^4 cm^{-3} or more (see e.g. Dahmen et al. 1998; Longmore et al. 2012, 2013b). There is also evidence to suggest that these clouds have higher values of α_{vir} than local clouds (Kauffmann, Pillai & Zhang 2013; Johnston et al. 2014).

To investigate this scenario, we performed two additional simulations that probe this regime. We modelled two clouds with mass $10^5 M_{\odot}$ and initial number density $n_0 = 10^4 \text{ cm}^{-3}$, in which the hydrogen was taken to be initially in molecular form. One of the clouds

was initialized with a turbulent field resulting in a virial parameter $\alpha_{\text{vir}} = 0.5$ (i.e. virial equilibrium), while the other had an $\alpha_{\text{vir}} = 2$ (i.e. the cloud was gravitationally unbound). These are shown as the green circles in Fig. 4. We see that the X-factors for these clouds are now much closer to the standard Milky Way value, demonstrating that the physical properties of the clouds play a significant role in controlling the X-factor. We also see that, as expected, the X-factor is lower in the case where the amount of turbulent kinetic energy in the cloud is higher.

Another factor that strongly affects the mean X-factor derived from the cloud maps is the value used for deciding whether or not a pixel should be included in the averaging process, $W_{\text{CO, cut}}$. If we set $W_{\text{CO, cut}} = 0$, and hence include all of the lines of sight in our synthetic emission maps when determining $\langle N_{\text{H}_2} \rangle$ and $\langle W_{\text{CO}} \rangle$, then we find that X_{CO} increases by between a factor of 3 (for the low-mass clouds) and an order of magnitude (for the high-mass clouds) as we increase the SFR from SFR_0 to 100 SFR_0 . However, if we restrict our calculation of the X-factor to only those lines of sight with detectable CO emission, we find that the dependence of X_{CO} on the SFR becomes much weaker. For example, if we set $W_{\text{CO, cut}} = 3 \text{ K km s}^{-1}$, we find that the X-factor increases at most by a factor of 2 as we increase the SFR from SFR_0 to 100 SFR_0 . This suggests that the main effect driving the increase in X_{CO} with SFR in the $W_{\text{CO, cut}} = 0$ case is a large increase in the number of lines of sight that have molecular hydrogen but no detectable CO emission, or in other words that pass through ‘CO-dark’ molecular gas (Wolfe et al. 2010). In those regions of the clouds that remain CO-bright, the relationship between CO emission and H_2 column density varies only weakly with the SFR. This behaviour is similar to what is seen in observational studies of the dependence of X_{CO} on metallicity: observations that focus only on CO-bright regions find a weak dependence of X_{CO} on Z (e.g. Wilson 1995; Bolatto et al. 2008), while those that sample both CO-bright and CO-dark gas find a much stronger dependence (e.g. Israel 1997; Bolatto et al. 2011). Finally, it is also interesting to note that the X-factor for the high-density, 100 SFR_0 clouds is insensitive to the value of $W_{\text{CO, cut}}$. This is because these clouds are so dense that the ‘skin’ in A_V -space that contains H_2 but not CO is so thin that very little mass is contained within it. As such, the CO typically always does a good job of tracing the molecular content of these clouds.

5 DISCUSSION

5.1 Consequences for the Kennicutt–Schmidt relation

The results presented in the previous section demonstrate that if we restrict our attention to those portions of molecular clouds that are traced by relatively bright CO emission, then the X-factor that we have to adopt in order to correctly convert from the CO intensity of these regions to their H_2 column density does not vary strongly with the local environment. Increasing the size of the assumed SFR (and hence the strength of the ISRF and the CRIR) by a factor of 100 leads to an increase in X_{CO} of around a factor of 2 at most. This result is consistent with the finding by Pineda et al. (2009) and Hughes et al. (2010) that in the Large Magellanic Cloud X_{CO} does not vary strongly with the strength of the ISRF, since these authors determine the masses of the clouds that they study using virial mass estimates that are sensitive only to the properties of the CO-bright gas.

Unfortunately, the number of galaxies for which we can distinguish between CO-bright and CO-faint portions of GMCs is very small. This is routinely done in the Milky Way, has already been

done to some extent in the Magellanic Clouds, and with the advent of the Atacama Large Millimeter/submillimeter Array (ALMA) becomes possible to do in more distant members of the Local Group. However, in most extragalactic observations, the best spatial resolution that we can achieve is comparable to or larger than the size of the individual clouds (see e.g. Colombo et al. 2014). In this case, we cannot easily distinguish between CO-bright and CO-faint portions of the clouds, and so the results that we present for the case when $W_{\text{CO, cut}} = 0$ give the best guide to the behaviour of X_{CO} in these systems.

For these unresolved clouds, our results show that X_{CO} increases substantially as we increase the SFR, provided that the other properties of the clouds remain unchanged (i.e. provided that their characteristic densities or velocity dispersions do not also vary as functions of the SFR). One important implication of this is that in order to be able to use CO emission as a reliable tracer of molecular mass, we need to understand the nature of the underlying cloud population and how this varies as a function of the SFR.

If we make the simplifying assumption that cloud properties do not depend to any great degree on the local value of the SFR, then our results imply that the standard interpretation of the Kennicutt–Schmidt relation may need to be reassessed. If the observed CO luminosity is dominated by high-mass clouds that have number densities similar to those of our fiducial cloud models, then our simulations suggest that X_{CO} should scale with the ambient rate of star formation as roughly³

$$X_{\text{CO}} \propto \text{SFR}^{1/2}. \quad (8)$$

Provided that the SFR does not vary widely over the area of the observational beam, it then follows that

$$X_{\text{CO}} \propto \Sigma_{\text{SFR}}^{1/2}. \quad (9)$$

Measurements of the Kennicutt–Schmidt relation at gas surface densities where the molecular component dominates (e.g. Bigiel et al. 2008; Blanc et al. 2009; Bigiel et al. 2011; Liu et al. 2011) find a power-law relationship that we can write as

$$\Sigma_{\text{SFR}} \propto \Sigma_{\text{mol}}^{N_{\text{obs}}}, \quad (10)$$

where N_{obs} denotes the observed power-law index. In practice, however, Σ_{mol} is usually not measured directly; instead, it is inferred from the CO luminosity through the application of some constant CO-to- H_2 conversion factor, which in normal spiral galaxies is generally taken to have the canonical Galactic value. Therefore, we can more accurately write the above relationship as

$$\Sigma_{\text{SFR}} \propto (\Sigma_{\text{CO, em}} X_{\text{CO, gal}})^{N_{\text{obs}}}, \quad (11)$$

where $\Sigma_{\text{CO, em}}$ is the surface density of CO emission. Now, if the actual CO-to- H_2 conversion factor is not fixed, but instead varies with Σ_{SFR} according to equation (9), then it follows that the actual relationship between Σ_{SFR} and Σ_{mol} can be written as

$$\Sigma_{\text{SFR}} \propto \Sigma_{\text{mol}}^{N_{\text{obs}}} \Sigma_{\text{SFR}}^{-N_{\text{obs}}/2}. \quad (12)$$

Rearranging this, we find that after correcting for the dependence of X_{CO} on the SFR, the true relationship between Σ_{SFR} and Σ_{mol} can be written as

$$\Sigma_{\text{SFR}} \propto \Sigma_{\text{mol}}^{N_{\text{act}}}, \quad (13)$$

³ Note that this scaling is approximate: it is not a precise fit to the simulation results.

where

$$N_{\text{act}} = \frac{2 N_{\text{obs}}}{2 + N_{\text{obs}}}. \quad (14)$$

Recent work by Bigiel et al. (2008) has suggested a value of $N_{\text{obs}} = 1$ from a sample of nearby galaxies. In this case, the ‘true’ index of the Kennicutt–Schmidt relation would be around $N_{\text{act}} = 2/3$, implying that CO is a progressively worse tracer of the SFR as one moves to more extreme environments (see the discussion in Shetty, Clark & Klessen 2014b). However, more recently, there have also been claims that $N_{\text{obs}} < 1$ (i.e. that the Kennicutt–Schmidt relation is sub-linear), even when assuming a constant X-factor (e.g. Ford et al. 2013; Shetty, Kelly & Bigiel 2013; Shetty et al. 2014a), with the most likely value lying at around 0.76. In this case, we would predict a value of $N_{\text{act}} = 0.55$.

It is important to note that this conclusion depends strongly on our assumption that the cloud properties do not vary strongly as a function of the local SFR. If this is not the case, and clouds in regions with higher values of Σ_{SFR} are systematically denser and/or more turbulent than clouds in less actively star-forming regions, then it is plausible that the decrease in X_{CO} caused by the higher densities and turbulent velocities could be large enough to overwhelm the increase caused by the higher SFR. Indeed, there is evidence that this is the case in the centres of many spiral galaxies: measurements of X_{CO} often find that it decreases close to the centre (see e.g. Sandstrom et al. 2013), while we know from the study of the Central Molecular Zone of our own galaxy that molecular clouds in this environment are very dense and highly turbulent (see e.g. the summary in the review of Molinari et al. 2014). Similarly, observational determinations of X_{CO} in nearby ultraluminous infrared galaxies (ULIRGs) or rapidly star-forming systems at high redshift typically find values that are smaller than the canonical local value (see e.g. Downes & Solomon 1998; Magdis et al. 2011; Ostriker & Shetty 2011; Hodge et al. 2012; Fu et al. 2013 or the recent review by Casey, Narayanan & Cooray 2014). However, these observations probe regions with much higher gas surface densities and turbulent velocities than are found in the local ISM, once again suggesting that the decrease in X_{CO} caused by the increased densities and velocities is able to overcome the increase caused by the higher radiation field strength and CRIR.

Ultimately, therefore, the real message to take away from this study is that it is necessary to understand the nature of the cloud population *before* one can properly interpret the CO emission. Unless one understands how the properties of the clouds depend on the local SFR, it is impossible to be certain whether X_{CO} increases with increasing Σ_{SFR} (as is the case in our simulations at fixed α_{vir}), decreases with increasing Σ_{SFR} (as appears to be the case in the centres of many spirals), or remains approximately constant. Further compounding this problem is the fact that CO measurements are one of the main observational tools used for determining the mass of molecular gas in clouds, potentially leaving us with a circular argument. Supplementary information on cloud masses, from tracers such as dust emission, can probably help to break this circularity, but a detailed examination of how best to do this is outside of the scope of this paper.

5.2 Comparison with previous work

It is interesting to compare our results on the dependence of the X-factor on the SFR with those of previous numerical studies that have looked at this problem. Bell et al. (2006) use the time-dependent UCL_PDR code (Papadopoulos, Thi & Viti 2002; Bell et al.

2005) to examine how X_{CO} varies as a function of depth within a semi-infinite slab of gas as they vary a number of different physical parameters, including the radiation field strength. They find that at high A_V , increasing the radiation field strength has little effect on X_{CO} : changing G_0 by a factor of 10^6 leads to no more than a factor of 3 change in X_{CO} . At low A_V , on the other hand, the dependence of X_{CO} on the radiation field is much stronger. At very low A_V , the gas is CO-faint and X_{CO} is very large, but as we move to higher A_V , the CO content of the gas grows, and the value of X_{CO} approaches the canonical Galactic value. The depth into the cloud at which this transition occurs depends on the radiation field strength, with stronger fields implying that the transition occurs at higher A_V . It is not straightforward to convert from these results to a cloud-averaged X_{CO} , but it is clear that they are at least qualitatively consistent with our findings that X_{CO} measured only for CO-bright regions does not vary strongly with the SFR, while X_{CO} averaged over the whole cloud (including the CO-faint regions) shows a much stronger dependence.

Feldmann et al. (2012) have also explored the dependence of X_{CO} on the local radiation field. They performed a large hydrodynamical simulation of the formation of a typical L_* galaxy (see also Gnedin & Kravtsov 2011) and then post-processed the results of this simulation using a sub-grid prescription for X_{CO} that assumes that it is primarily determined by the mean visual extinction of the individual clouds and the strength of the ISRF. For $G_0 = 1$, Feldmann et al. (2012) use the results of Glover & Mac Low (2011) to calibrate their sub-grid model, while for $G_0 \gg 1$ and $G_0 \ll 1$, they estimate the CO abundance (and hence X_{CO}) using a simple remapping procedure that assumes that the clouds are in photodissociation equilibrium. They find that for clouds with mean extinctions $\bar{A}_V \sim 6$ and below, the value of X_{CO} depends strongly on the radiation field strength. For higher column density clouds, on the other hand, X_{CO} becomes largely independent of the radiation field strength. As most of the molecular clouds that form in their simulations have relatively large column densities, they find that when you average over the whole population of clouds, the dependence of X_{CO} on the strength of the ISRF becomes very weak.

At first sight, these results would appear to contradict our finding that X_{CO} can vary significantly as we change the SFR. However, there are a couple of important points that one should bear in mind. First, Feldmann et al. (2012) do not vary the CRIR, only the radiation field strength. Therefore, their models do not account for the increased destruction of CO by dissociative charge transfer with He^+ that occurs when the CRIR is large. At high A_V , this is the main CO destruction mechanism, and its inclusion in our models is one of the main reasons why we see at least some dependence of X_{CO} on SFR even when we restrict our attention to portions of the cloud that are highly shielded and CO-bright. Secondly, the hydrodynamical simulations used as a basis for the Feldmann et al. (2012) analysis have a resolution of only 60 pc and hence are sensitive only to the largest, densest clouds, which are naturally the ones least affected by changes in the SFR.

Another large-scale numerical study of the dependence of X_{CO} on environment was presented by Narayanan et al. (2012). They post-processed an extensive series of SPH simulations of isolated and merging galaxies in order to determine the local chemical and thermal state of the gas and the consequent CO emission (see also Krumholz, Leroy & McKee 2011; Narayanan et al. 2011). They found that in their models, X_{CO} scales inversely with the molecular gas surface density as $X_{\text{CO}} \propto \Sigma_{\text{H}_2}^{-1/2}$. In these models, $\Sigma_{\text{SFR}} \propto \Sigma_{\text{H}_2}^{3/2}$ by construction, and so the implication is that X_{CO} should decrease weakly with increasing SFR, scaling as $X_{\text{CO}} \propto \Sigma_{\text{SFR}}^{-1/3}$. However, as

with many of the observational studies of X_{CO} in extreme environments mentioned in the previous section, direct comparison of these results with our own is complicated by the fact that the mean properties of the clouds in regions with high SFR in these simulations are not the same as those of the clouds in regions with low SFR.

Lagos et al. (2012) also studied the effects of changes in the strength of the ISRF and the CRIR on the value of X_{CO} using a simple semi-analytical approach. They used the `GALFORM` semi-analytical galaxy formation model (Lagos et al. 2011a,b) to compute the atomic and molecular gas content of a series of model galaxies, and then used the `UCL_PDR` code to determine the CO luminosity of the molecular gas, from which they could then infer X_{CO} . In their photon dominated region (PDR) calculations, it was necessary for them to assume some representative number density and visual extinction for the molecular gas. In most of their models, Lagos et al. (2012) assumed a characteristic number density $n = 10^4 \text{ cm}^{-3}$ and a characteristic extinction $A_V = 8$. In these conditions, one would expect CO photodissociation to be completely negligible even when the ISRF is very strong, meaning that changes in the radiation field strength will only affect X_{CO} indirectly, through its effect on the thermal balance of the dust. Indeed, Lagos et al. (2012) find that in their model, even very large changes in the radiation field strength have only a small influence on X_{CO} . However, in practice, in realistic GMC models only a small fraction of the total gas mass is found in regions with a mean visual extinction as high as $A_V = 8$ (see e.g. Clark & Glover 2014), and so by focusing only on these conditions, Lagos et al. (2012) ignore the large changes in the CO content and CO luminosity of the lower density, less shielded gas that drive much of the variation that we find in X_{CO} . Lagos et al. (2012) also examine the effect of varying the CRIR while keeping the radiation field strength fixed, and find that in this case, X_{CO} increases slowly with increasing ionization rate, in general agreement with our results.

The dependence of X_{CO} on the UV field strength was also examined by Offner et al. (2014). They computed X_{CO} for clouds illuminated by radiation fields with 1 and 10 times the strength of the Draine (1978) field, using the `3D-PDR` code (Bisbas et al. 2012). When computing X_{CO} , they only considered pixels in their synthetic emission maps with CO integrated intensities $W_{\text{CO}} > 0.45 \text{ K km s}^{-1}$. They find that $X_{\text{CO}} = 1.5X_{\text{CO,gal}}$ for their $G_0 = 1.7$ run and $X_{\text{CO}} = 2.15X_{\text{CO,gal}}$ for their $G_0 = 17$ run, a roughly 50 per cent increase. This is somewhat smaller than the difference we find between our $\text{SFR} = \text{SFR}_0$ and $\text{SFR} = 10\text{SFR}_0$ runs, which for a similar integrated intensity cut show values of X_{CO} that differ by closer to a factor of 2. However, it is important to note that Offner et al. (2014) do not vary the CRIR at the same time that they vary the radiation field strength, and so it is likely that there is less destruction of CO in their $G_0 = 17$ run than in our $\text{SFR} = 10\text{SFR}_0$ run.

Finally, in a recent paper, Bisbas, Papadopoulos & Viti (2015) have studied how the typical CO abundance in a GMC varies as a function of the CRIR. They find that increasing the CRIR by a factor of 10 or more leads to substantial destruction of CO and argue that this will lead to clouds in regions with $\text{SFR} > 10\text{SFR}_0$ being CO-faint. However, they do not compute X_{CO} for any of their clouds models and hence do not quantify its dependence on the SFR.

5.3 Caveats

There are two important methodological caveats that the reader should bear in mind when considering our results. First, as we have already discussed in Section 4.2, we do not currently account for

the effects of line-of-sight velocity gradients when determining the effectiveness of H_2 self-shielding. This means that we will tend to overestimate the effectiveness of self-shielding in gas with H_2 column densities of around $10^{14} < N_{\text{H}_2} < 10^{18} \text{ cm}^{-2}$, the regime where Doppler broadening of the Lyman–Werner lines dominates the UV absorption spectrum. Consequently, we will overestimate the total H_2 content of the cloud. However, as most of the H_2 in the cloud is found in regions with much higher H_2 column densities, we would not expect the omission of this effect to have a large impact on our results.

Secondly, we have assumed that the CRIR is uniform throughout the cloud, or in other words that the cosmic ray absorption spectrum is not significantly affected by absorption within the cloud. Whether or not this is a good approximation depends on the details of the low-energy portion of the cosmic ray energy spectrum, which is poorly constrained (Padovani, Galli & Glassgold 2009). There is some tentative evidence from astrochemical studies that the CRIR in dense cores may be significantly lower than in the diffuse ISM (see e.g. Indriolo & McCall 2012), but the issue is far from being settled. If there is indeed a significant fall-off in the CRIR as we move from diffuse to dense gas, then our models will tend to overestimate the heating rate in the dense gas, and also the rate at which CO is destroyed there. The net effect of this on X_{CO} is difficult to predict without detailed modelling. However, we note that the destruction of the relatively diffuse interclump CO component that occurs as we increase the SFR is a consequence of the increasing strength of the ISRF, and not the increase in the CRIR. Therefore, even if cosmic ray absorption within the cloud becomes significant, this behaviour will remain the same, and so we would expect to still see an increase in X_{CO} as we increase the SFR.

In addition to these methodological caveats, we also take this opportunity to remind the reader that when we refer to changes in the ‘SFR’, what we mean are changes in the strength of the ISRF and the size of the CRIR, since we assume that both scale linearly with the SFR. It could well be that this assumption is too simplistic, and that the functional dependence of the radiation field strength and the CRIR on the local SFR is more complicated than a simple linear scaling. If so, then the results of our study will still hold, but the actual SFRs corresponding to our 10SFR_0 and 100SFR_0 models will differ from what we have assumed here.

6 CONCLUSIONS

In this paper, we have performed a series of numerical simulations of molecular cloud evolution that explore how the distribution of CO emission coming from the clouds changes as we change the strength of the ISRF and the size of the CRIR in the clouds. If one makes the reasonable assumption that these quantities scale linearly with the local SFR, then our results indicate how the CO-to- H_2 conversion factor for the clouds, X_{CO} , depends on the SFR.

We find that as we increase the radiation field strength and the CRIR, the CO content of our simulated clouds decreases. CO survives more effectively in well-shielded clump and filaments than in the more diffuse interclump gas, leading to a significant decrease in the filling factor of bright CO emission. The integrated intensity of the CO emission from the brightest regions increases with increasing SFR, owing to the heating of the gas by the higher cosmic ray flux, but this increase is too small to compensate for the loss of CO emission from elsewhere in the cloud, and so overall the CO luminosity of the cloud decreases.

The change in the chemical composition of the cloud as we increase the SFR leads to a change in X_{CO} . However, the size of

this change depends on the method we use to compute X_{CO} . If we consider the whole area of the cloud when computing the mean H_2 column density and CO integrated intensity, then we find that X_{CO} changes substantially, scaling as roughly $X_{\text{CO}} \propto \text{SFR}^{1/2}$ for our high-mass clouds. On the other hand, if we consider only regions with CO integrated intensities exceeding an intensity threshold $W_{\text{CO, cut}}$, then we find that the dependence of X_{CO} on SFR weakens substantially as we increase $W_{\text{CO, cut}}$. The reason for this difference in behaviour is that the fraction of the cloud filled with ‘dark’ molecular gas (i.e. H_2 without associated CO) increases as we increase the SFR. This dark gas contributes to our calculation of X_{CO} in the case where we do not apply an intensity threshold, but is ignored when we do apply a threshold.

We have also explored whether increasing the density and turbulent velocity dispersion of our model clouds at the same time as we increase the SFR affects the relationship between X_{CO} and the SFR. We find that if we increase the initial number density to $n_0 = 10^4 \text{ cm}^{-3}$ and raise the virial parameter to $\alpha_{\text{vir}} = 2$, then we can recover a value of X_{CO} close to the canonical Galactic value even when the local SFR is 100 times our default value. We conclude from this that we will only recover a positive correlation between X_{CO} and the SFR if there is not a strong correlation between mean cloud densities and the SFR. The fact that observational determinations of X_{CO} in starbursting systems such as ULIRGs find a value *lower* than the canonical Galactic value therefore provides a strong indication that the properties of typical molecular clouds in these systems must differ significantly from those of molecular clouds in the local ISM. Consequently, we argue that the most important message to take away from this study is that it is necessary to understand the nature of the cloud population before one can properly interpret the CO emission coming from the clouds.

Finally, we stress that if the CO emission from nearby galaxies is dominated by virialized clouds with masses around $10^5 M_\odot$ and densities around a few 100 cm^{-3} , then the X-factor may vary significantly as the local SFR changes. In these circumstances, our standard interpretation of the Kennicutt–Schmidt relation, in which X_{CO} is assumed to be constant throughout the galaxy, should be rethought.

ACKNOWLEDGEMENTS

The authors would like to thank J. Pety, R. Feldmann, N. Gnedin and F. Israel for interesting discussions concerning the physics of the CO-to- H_2 conversion factor in strong radiation fields. SCOG acknowledges financial support from the Deutsche Forschungsgemeinschaft via SFB 881, ‘The Milky Way System’ (sub-projects B1, B2 and B8) and SPP 1573, ‘Physics of the Interstellar Medium’ (grant number GL 668/2-1).

REFERENCES

Ballesteros-Paredes J., Mac Low M.-M., 2002, *ApJ*, 570, 734
 Ballesteros-Paredes J., Hartmann L. W., Vázquez-Semadeni E., Heitsch F., Zamora-Avilés M. A., 2011, *MNRAS*, 411, 65
 Bate M. R., Burkert A., 1997, *MNRAS*, 288, 1060
 Beaumont C. N., Offner S. S. R., Shetty R., Glover S. C. O., Goodman A. A., 2013, *ApJ*, 777, 173
 Bell T. A., Viti S., Williams D. A., Crawford I. A., Price R. J., 2005, *MNRAS*, 357, 961
 Bell T. A., Roueff E., Viti S., Williams D. A., 2006, *MNRAS*, 371, 1865
 Bergin E. A., Tafalla M., 2007, *ARA&A*, 45, 339
 Bigiel F., Leroy A., Walter F., Brinks E., de Blok W. J. G., Madore B., Thornley M. D., 2008, *AJ*, 136, 2846

Bigiel F. et al., 2011, *ApJ*, 730, L13
 Bisbas T. G., Bell T. A., Viti S., Yates J., Barlow M. J., 2012, *MNRAS*, 427, 2100
 Bisbas T. G., Papadopoulos P. P., Viti S., 2015, *ApJ*, 803, 37
 Black J. H., 1994, in Cutri R. M., Latter W. B., eds, *ASP Conf. Ser. Vol. 58, The First Symposium on the Infrared Cirrus and Diffuse Interstellar Clouds*. Astron. Soc. Pac., San Francisco, p. 355
 Blanc G. A., Heiderman A., Gebhardt K., Evans N. J., Adams J., 2009, *ApJ*, 704, 842
 Blitz L., Fukui Y., Kawamura A., Leroy A., Mizuno N., Rosolowsky E., 2007, in Reipurth B., Jewitt D., Keil K., eds, *Protostars & Planets V*. Univ. Arizona Press, Tucson, p. 81
 Bolatto A. D., Leroy A. K., Rosolowsky E., Walter F., Blitz L., 2008, *ApJ*, 686, 948
 Bolatto A. D. et al., 2011, *ApJ*, 741, 12
 Bolatto A. D., Wolfire M., Leroy A. K., 2013, *ARA&A*, 51, 207
 Bonatto C., Bica E., 2011, *MNRAS*, 415, 2827
 Casey C. M., Narayanan D., Cooray A., 2014, *Phys. Rep.*, 541, 45
 Clark P. C., Glover S. C. O., 2014, *MNRAS*, 444, 2396
 Clark P. C., Glover S. C. O., Klessen R. S., 2012a, *MNRAS*, 420, 745
 Clark P. C., Glover S. C. O., Klessen R. S., Bonnell I. A., 2012b, *MNRAS*, 424, 2599
 Clark P. C., Glover S. C. O., Ragan S. E., Shetty R., Klessen R. S., 2013, *ApJ*, 768, L34
 Colombo D. et al., 2014, *ApJ*, 784, 3
 Dahmen G., Huttemeister S., Wilson T. L., Mauersberger R., 1998, *A&A*, 331, 959
 Dame T. M., Hartmann D., Thaddeus P., 2001, *ApJ*, 547, 792
 Dickman R. L., 1978, *ApJS*, 37, 407
 Digel S. W., Aprile E., Hunter S. D., Mukherjee R., Xu F., 1999, *ApJ*, 520, 196
 Dobbs C. L., Burkert A., Pringle J. E., 2011, *MNRAS*, 413, 2935
 Downes D., Solomon P. M., 1998, *ApJ*, 507, 615
 Draine B. T., 1978, *ApJS*, 36, 595
 Elmegree B. G., Klessen R. S., Wilson C. D., 2008, *ApJ*, 681, 365
 Feldmann R., Gnedin N. Y., Kravtsov A. V., 2012, *ApJ*, 747, 124
 Ferriere K. M., 2001, *Rev. Mod. Phys.*, 73, 1031
 Ford G. P. et al., 2013, *ApJ*, 769, 55
 Froebrich D., Rowles J., 2010, *MNRAS*, 406, 1350
 Fu H. et al., 2013, *Nature*, 498, 338
 Glover S. C. O., Clark P. C., 2012a, *MNRAS*, 421, 116
 Glover S. C. O., Clark P. C., 2012b, *MNRAS*, 426, 377
 Glover S. C. O., Mac Low M.-M., 2007a, *ApJS*, 169, 239
 Glover S. C. O., Mac Low M.-M., 2007b, *ApJ*, 659, 1317
 Glover S. C. O., Mac Low M.-M., 2011, *MNRAS*, 412, 337
 Glover S. C. O., Federrath C., Mac Low M.-M., Klessen R. S., 2010, *MNRAS*, 404, 2
 Gnedin N. Y., Kravtsov A. V., 2011, *ApJ*, 728, 88
 Goldsmith P. F., 2001, *ApJ*, 557, 736
 Goodman A. A., Pineda J. E., Schnee S. L., 2009, *ApJ*, 692, 91
 Gredel R., Lepp S., Dalgarno A., 1987, *ApJ*, 323, L137
 Habing H. J., 1968, *Bull. Astron. Inst. Netherlands*, 19, 421
 Heyer M., Krawczyk C., Duval J., Jackson J. M., 2009, *ApJ*, 699, 1092
 Hodge J. A., Carilli C. L., Walter F., de Blok W. J. G., Riechers D., Daddi E., Lentati L., 2012, *ApJ*, 760, 11
 Hughes A. et al., 2010, *MNRAS*, 406, 2065
 Indriolo N., McCall B. J., 2012, *ApJ*, 745, 91
 Israel F. P., 1997, *A&A*, 328, 471
 Johnston K. G., Beuther H., Linz H., Schmiedeke A., Ragan S. E., Henning Th., 2014, *A&A*, 568, A56
 Kainulainen J., Beuther H., Henning T., Plume R., 2009, *A&A*, 508, L35
 Kauffmann J., Pillai T., Zhang Q., 2013, *ApJ*, 765, L35
 Krumholz M. R., McKee C. F., Tumlinson J., 2008, *ApJ*, 689, 865
 Krumholz M. R., Leroy A. K., McKee C. F., 2011, *ApJ*, 731, 25
 Lagos C. D., Lacey C. G., Baugh C. M., Bower R. G., Benson A. J., 2011a, *MNRAS*, 416, 1566
 Lagos C. D., Baugh C. M., Lacey C. G., Benson A. J., Kim H.-S., Power C., 2011b, *MNRAS*, 418, 1649

- Lagos C. D. et al., 2012, MNRAS, 426, 2142
- Larson R. B., 1981, MNRAS, 194, 809
- Lee H.-H., Herbst E., Pineau des Forêts G., Roueff E., Le Bourlot J., 1996, A&A, 311, 690
- Lee M.-Y., Stanimirović S., Wolfire M. G., Shetty R., Glover S. C. O., Molina F. Z., Klessen R. S., 2014, ApJ, 784, 80
- Liu G., Koda J., Calzetti D., Fukuhara M., Momose R., 2011, ApJ, 735, 63
- Longmore S. et al., 2012, ApJ, 746, 117
- Longmore S. et al., 2013a, MNRAS, 429, 987
- Longmore S. et al., 2013b, MNRAS, 433, L15
- Magdis G. E. et al., 2011, ApJ, 740, L15
- Molina F., Glover S. C. O., Federrath C., 2011, in Röllig M., Simon R., Ossenkopf V., Stutzki J., eds, Conditions and Impact of Star Formation, EAS Publications Ser. Vol. 52, p. 289
- Molinari S. et al., 2014, in Beuther H., Klessen R. S., Dullemond C. P., Henning T., eds, Protostars and Planets VI. Univ. Arizona Press, Tucson, p. 125
- Narayanan D., Hopkins P. F., 2013, MNRAS, 433, 1223
- Narayanan D., Krumholz M. R., Ostriker E. C., Hernquist L., 2011, MNRAS, 418, 664
- Narayanan D., Krumholz M. R., Ostriker E. C., Hernquist L., 2012, MNRAS, 421, 3127
- Nelson R. P., Langer W. D., 1999, ApJ, 524, 923
- Offner S. S. R., Bisbas T. G., Bell T. A., Viti S., 2014, MNRAS, 440, L81
- Ostriker E. C., Shetty R., 2011, ApJ, 731, 41
- Padovani M., Galli D., Glassgold A. E., 2009, A&A, 501, 619
- Papadopoulos P. P., Thi W.-F., Viti S., 2002, ApJ, 579, 270
- Pineda J. E., Caselli P., Goodman A. A., 2008, ApJ, 679, 481
- Pineda J. L., Ott J., Klein U., Wong T., Muller E., Hughes A., 2009, ApJ, 703, 736
- Pineda J. L., Goldsmith P. F., Chapman N., Snell R. L., Li D., Cambresy L., Brunt C., 2010, ApJ, 721, 686
- Roman-Duval J., Jackson J. M., Heyer M., Rathborne J., Simon R., 2010, ApJ, 723, 492
- Sanders D. B., Solomon P. M., Scoville N. Z., 1984, ApJ, 276, 182
- Sandstrom K. M. et al., 2013, ApJ, 777, 5
- Schöier F. L., van der Tak F. F. S., van Dishoeck E. F., Black J. H., 2005, A&A, 432, 369
- Scoville N. Z., Yun M. S., Sanders D. B., Clemens D. P., Waller W. H., 1987, ApJS, 63, 821
- Sembach K. R., Howk J. C., Ryans R. S. I., Keenan F. P., 2000, ApJ, 528, 310
- Shetty R., Glover S. C., Dullemond C. P., Klessen R. S., 2011a, MNRAS, 412, 1686
- Shetty R., Glover S. C., Dullemond C. P., Ostriker E. C., Harris A. I., Klessen R. S., 2011b, MNRAS, 415, 3253
- Shetty R., Kelly B. C., Bigiel F., 2013, MNRAS, 430, 288
- Shetty R. et al., 2014a, MNRAS, 437, L61
- Shetty R., Clark P. C., Klessen R. S., 2014b, MNRAS, 442, 2208
- Smith R. J., Glover S. C. O., Clark P. C., Klessen R. S., Springel V., 2014, MNRAS, 441, 1628
- Sobolev V. V., 1957, SvA, 1678
- Solomon P. M., Rivolo A. R., Barrett J., Yahil A., 1987, ApJ, 319, 730
- Springel V., 2005, MNRAS, 364, 1105
- Sternberg A., Le Petit F., Roueff E., Le Bourlot J., 2014, ApJ, 790, 10
- Strong A. W., Mattox J. R., 1996, A&A, 308, L21
- Visser R., van Dishoeck E. F., Black J. H., 2009, A&A, 503, 323
- Ward R. L., Wadsley J., Sills A., 2014, MNRAS, 439, 651
- Wilson C. D., 1995, ApJ, 448, L97
- Wolfire M. G., Hollenbach D., McKee C. F., 2010, ApJ, 716, 1191
- Woodall J., Agúndez M., Markwick-Kemper A. J., Millar T. J., 2007, A&A, 466, 1197

This paper has been typeset from a \LaTeX file prepared by the author.